

BeSS report – September 2014

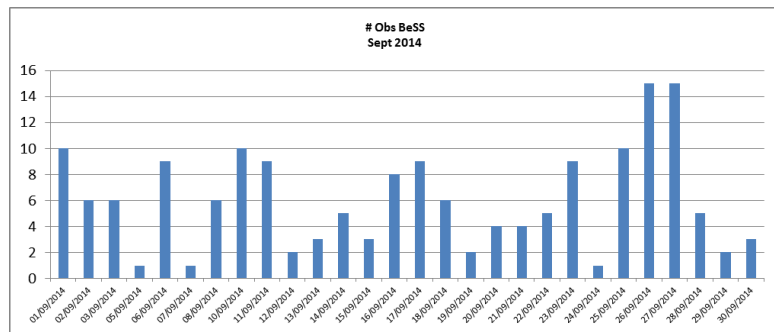
Data compiled by Valérie Desnoux

Do not miss the new section on the Be projects by Ernst Pollmann [here](#)

Observateur	Nb spec
Pujol	29
HOUPERT	22
Fosanelli	19
Graham	16
Buil	16
Sawicki	15
GARREL	14
Bohlsen	9
Lemoult	8
Pollmann	6
LAILLY	5
Berardi	4
MontigianiMannucci	3
Guarro Fló	3
Total général	169

- 169 H-alpha spectra acquired
- 109 objects observed
- 14 observers contributed

The most observed objects were gam Cas, pi Aqr and HD 194057

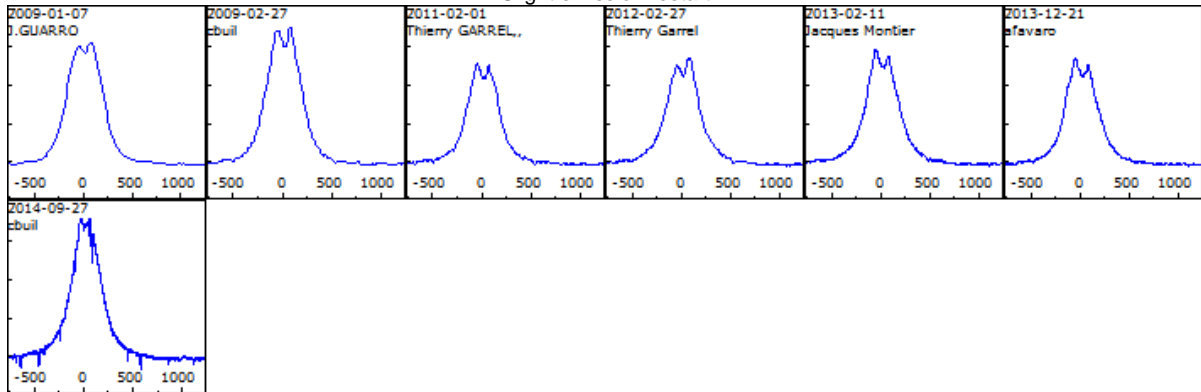


Objects observed

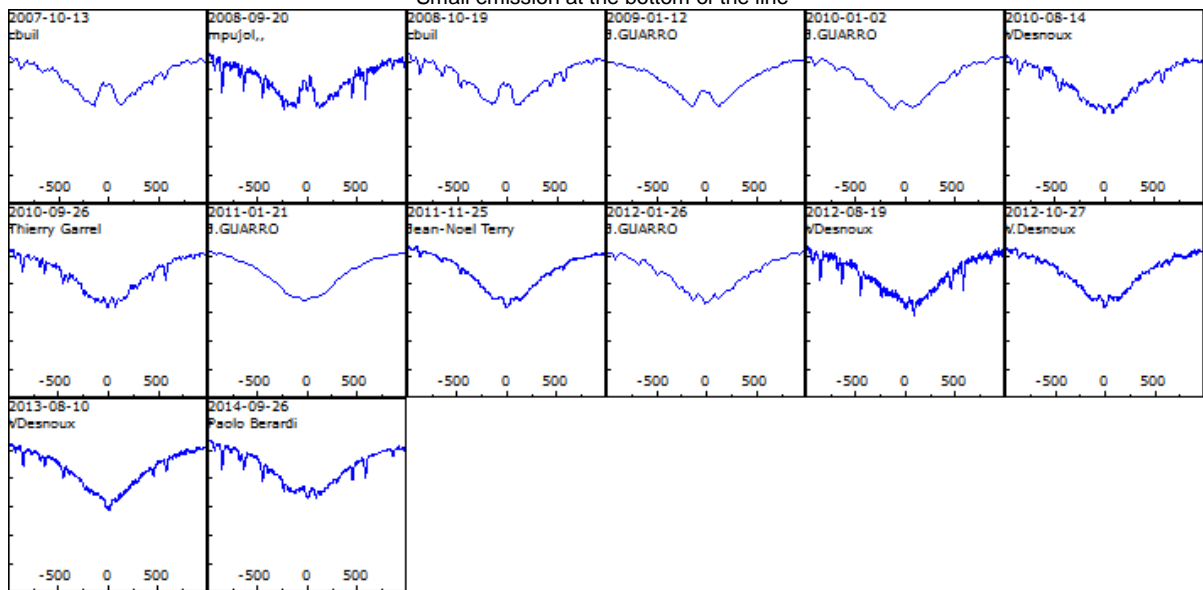
Classique							Herbig
gam Cas	ALCYONE	EM* VES 704	phi Per	13 Tau	HD 206773	V408 Lac	HD 31648
pi Aqr	iot Lyr	HD 153199	HD 157273	FR CMa	MEROPE	HD 208682	
HD 194057	phi And	tet CrB	60 Cyg	eps Cap	V848 Ara	V396 Vul	
lam Cyg	66 Oph	BD+56 251	lam Eri	HD 192445	HD 2789	HD 24479	
QR Vul	V2136 Cyg	59 Cyg BD+36	25 Cyg	BD+46 3087	HD 183339	V787 Cas	
del Sco	HD 196712	3956B	V1075 Sco	HD 153262	iot Ara	HD 201522	
HD 201836	nu Cyg	12 Vul	HD 195554	BD+56 259	chi Oph	V742 Cas	
V813 Cas	eps Cas	zet Tau	V2148 Cyg	V341 Sge	HD 13669	HD 179343	
omi Her	14 Lac	25 Vul	ELECTRA	V783 Cas	eta PsA	V846 Ara	
PLEIONE	25 Ori	BD+62 300	BD+62 245	HD 199218	V818 Cas	HD 169033	
V442 And	zet Oph	ups Sgr	28 Cyg	EM* AS 28	bet Psc	V2163 Cyg	
4 Aql	HD 9612	V782 Cas	V2139 Cyg	HD 20134	V2119 Cyg	LQ And	
V764 Cas	KX And	HD 193182	BD+60 2600	EM* AS 16	48 Lib	kap Aql	
31 Peg	omi Aqr	NT Peg	V423 Lac	HD 9709	V2153 Cyg	48 Per	
KY And	QV Tel	omi Cas	ome Ori	HD 228860	105 Tau	V568 Cyg	
HD 153879	V777 Cas	HD 194779					

Emission increase since last observations

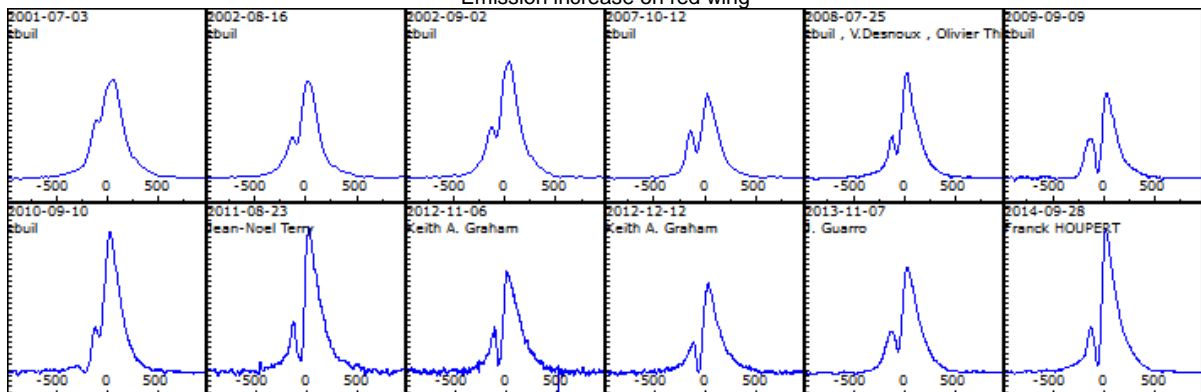
FR CMa
Slight emission restart



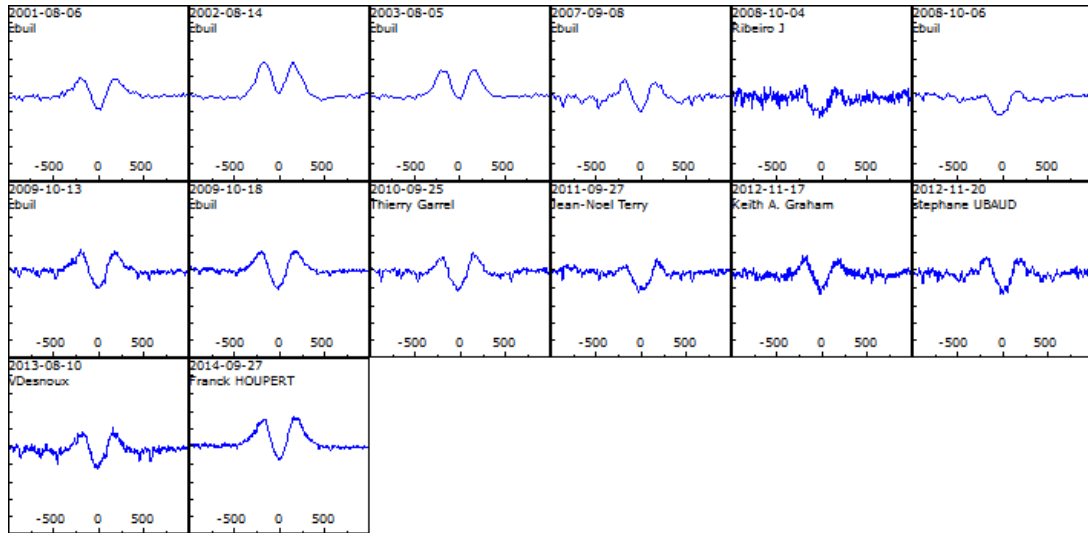
HD 24479
Small emission at the bottom of the line



KX And
Emission increase on red wing

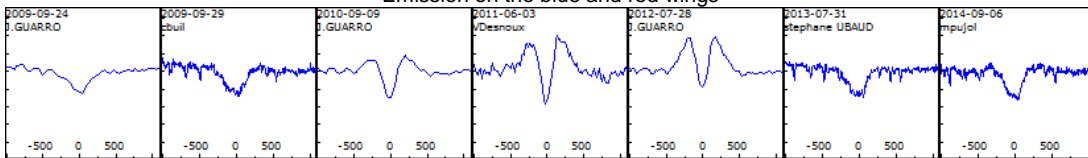


KY And
Emission restart as in 2002



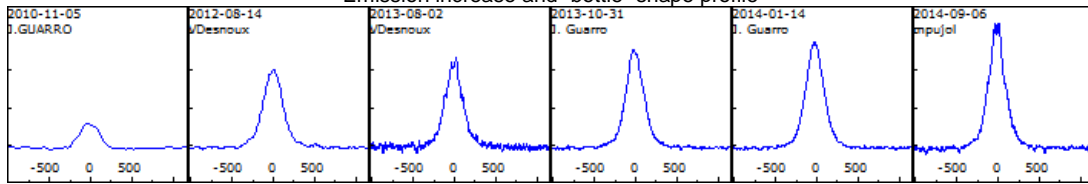
V396 Vul

Emission on the blue and red wings



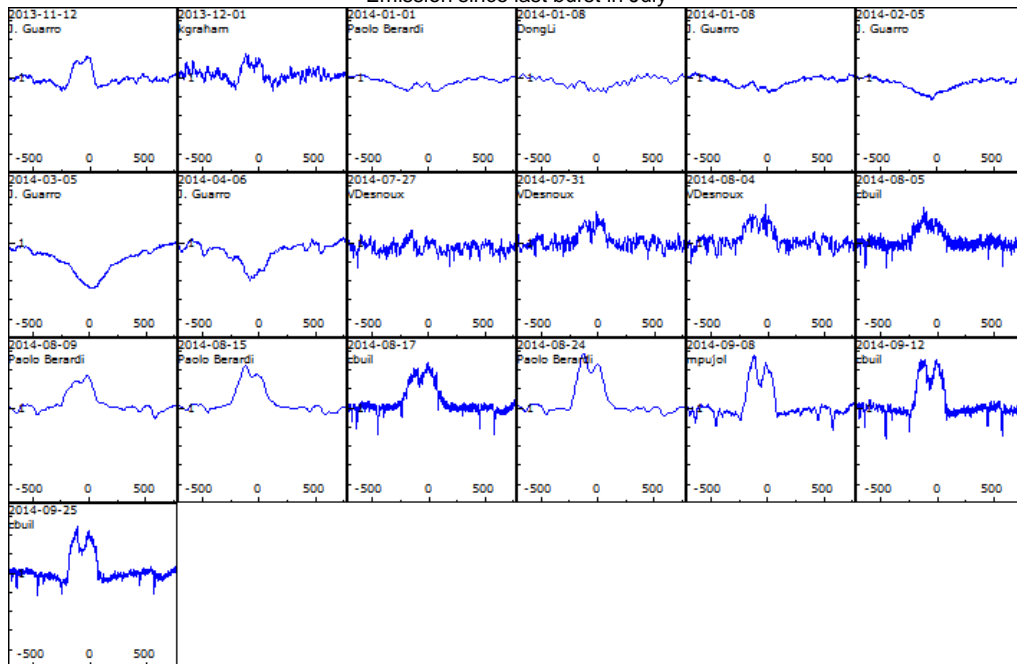
V408 Lac

Emission increase and "bottle" shape profile



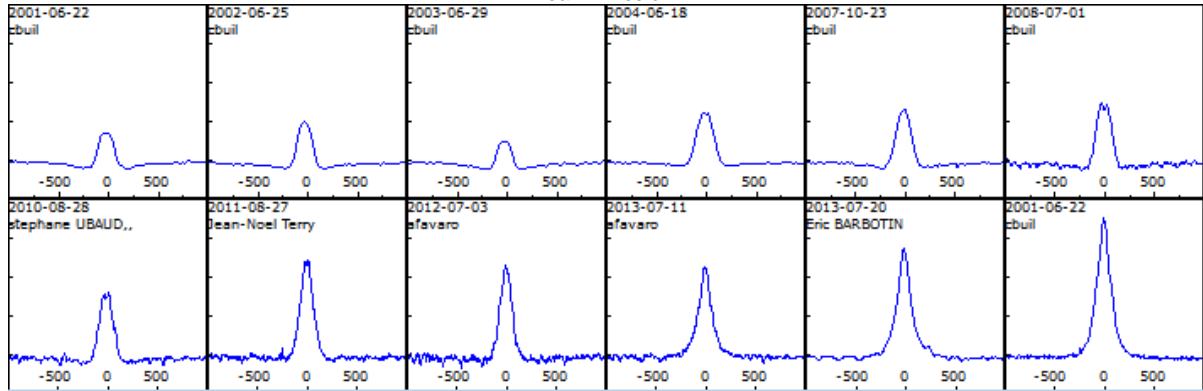
V442 And

Emission since last burst in July



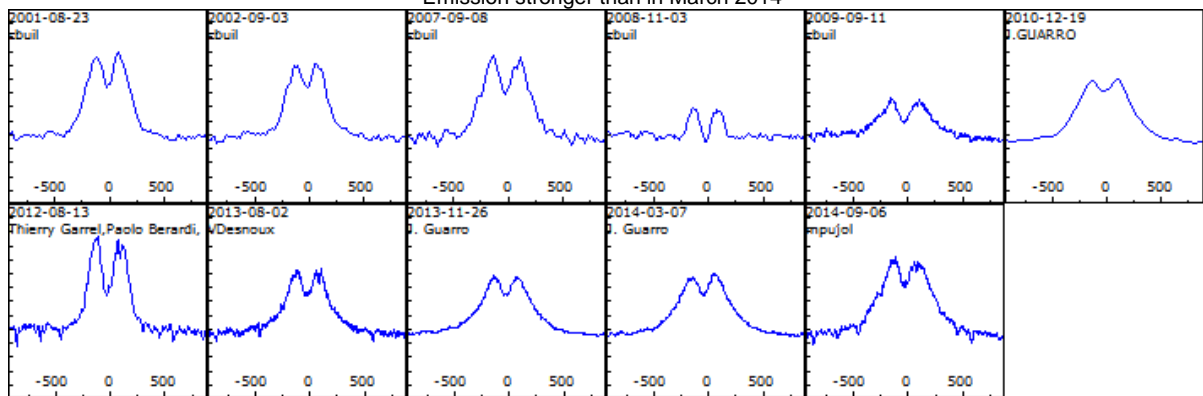
V586 Cyg

Peak Emission



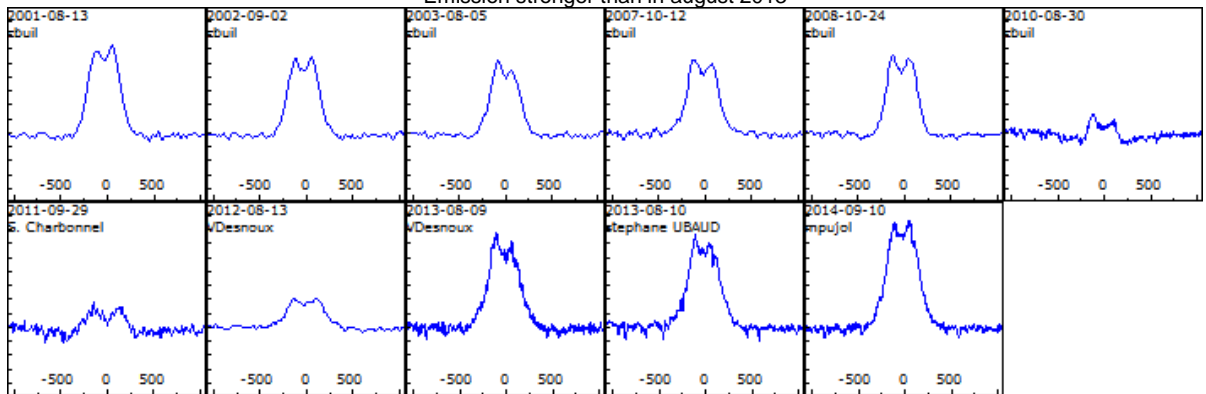
V787 cas

Emission stronger than in March 2014



V818 Cas

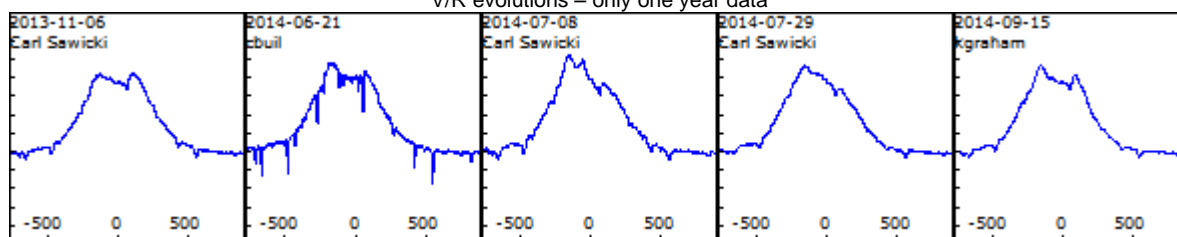
Emission stronger than in august 2013



Moderate evolutions of H-alpha line

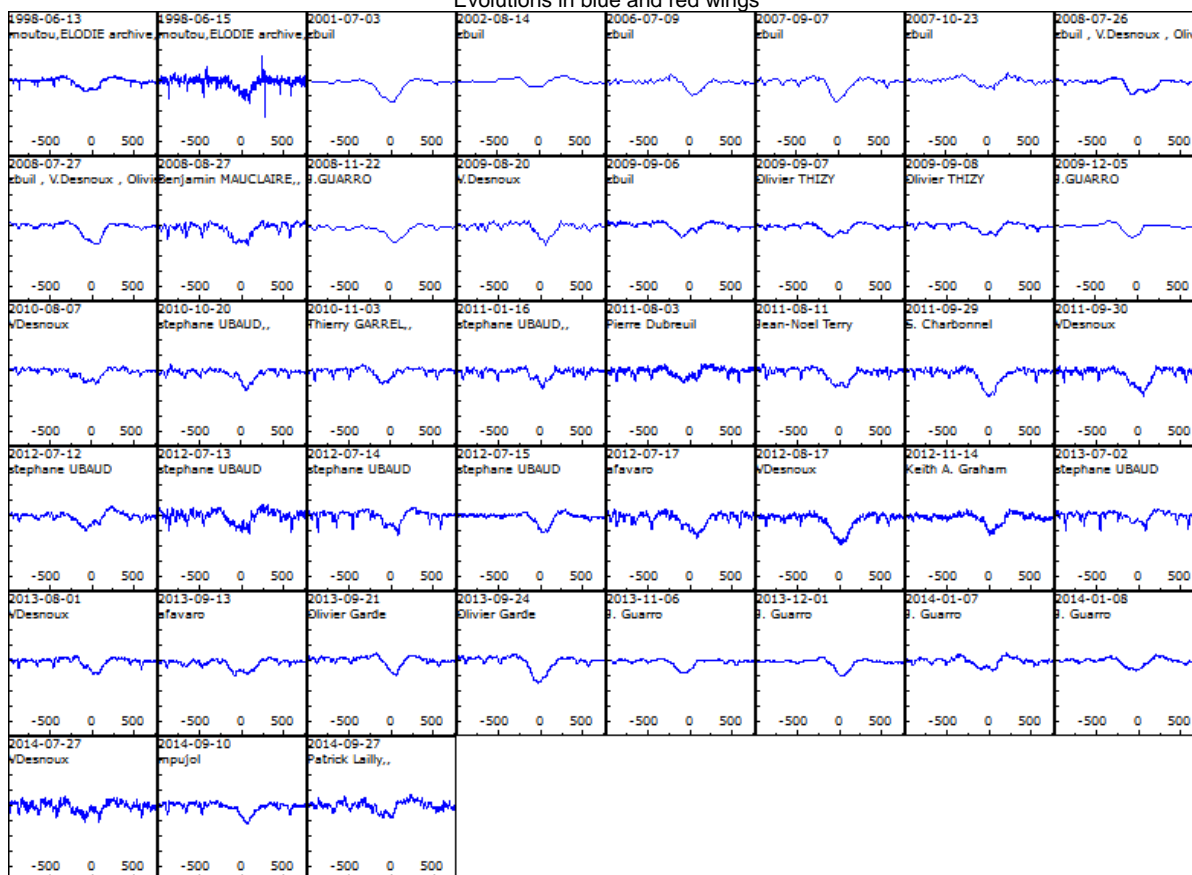
59 Cyg

V/R evolutions – only one year data



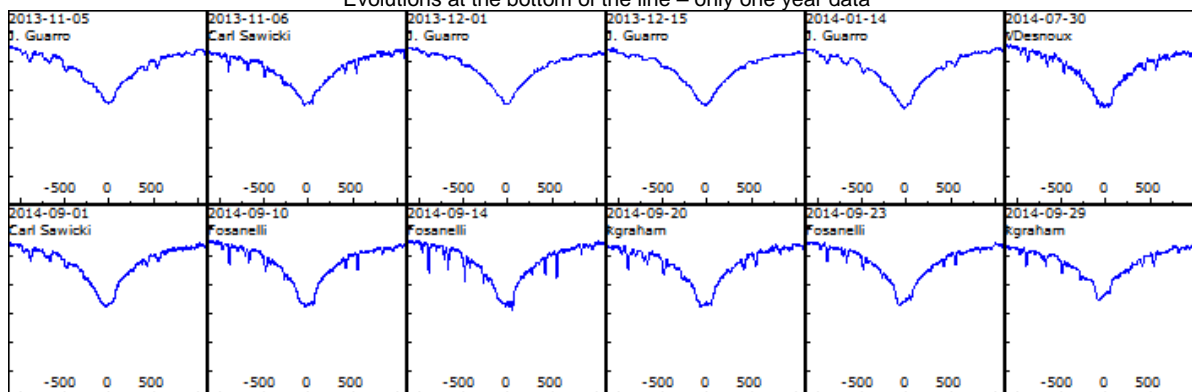
14 Lac

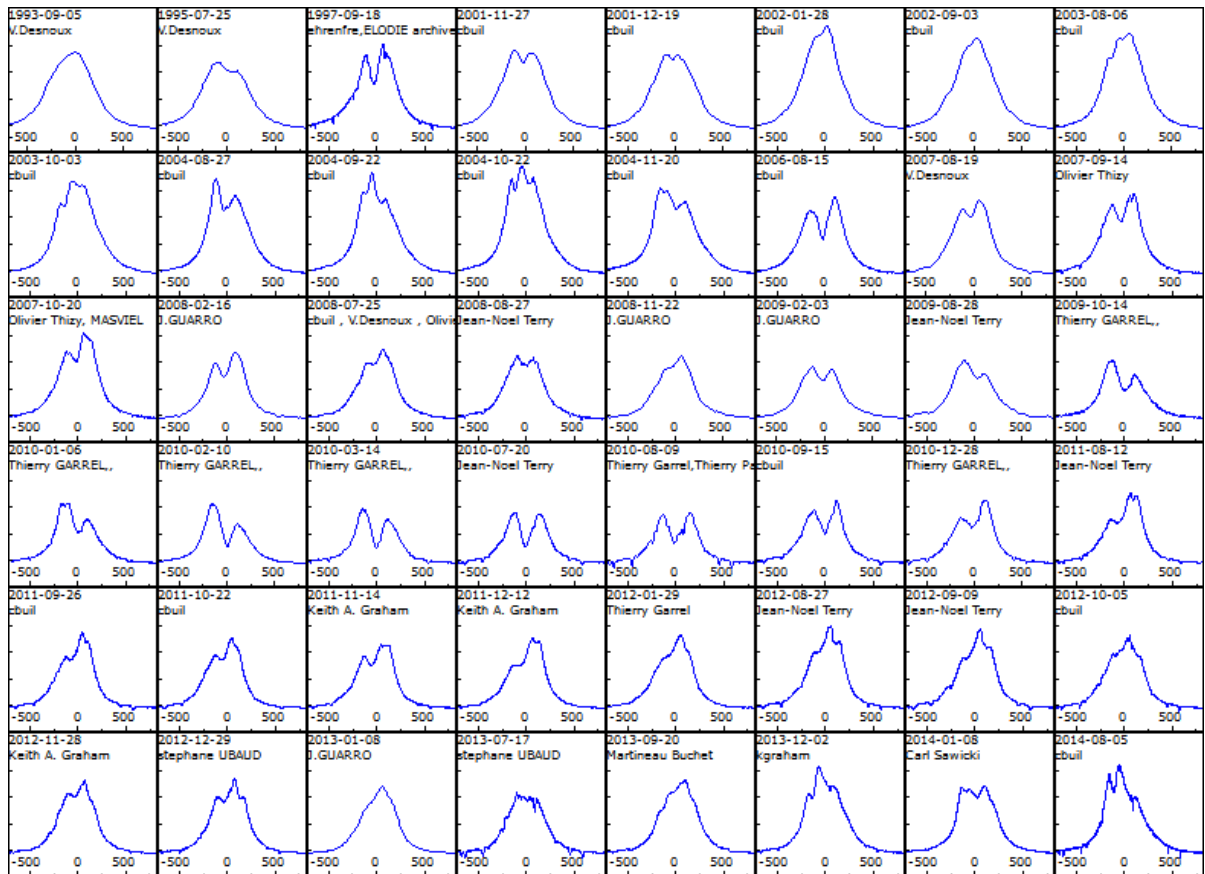
Evolution in blue and red wings



Lam cyg

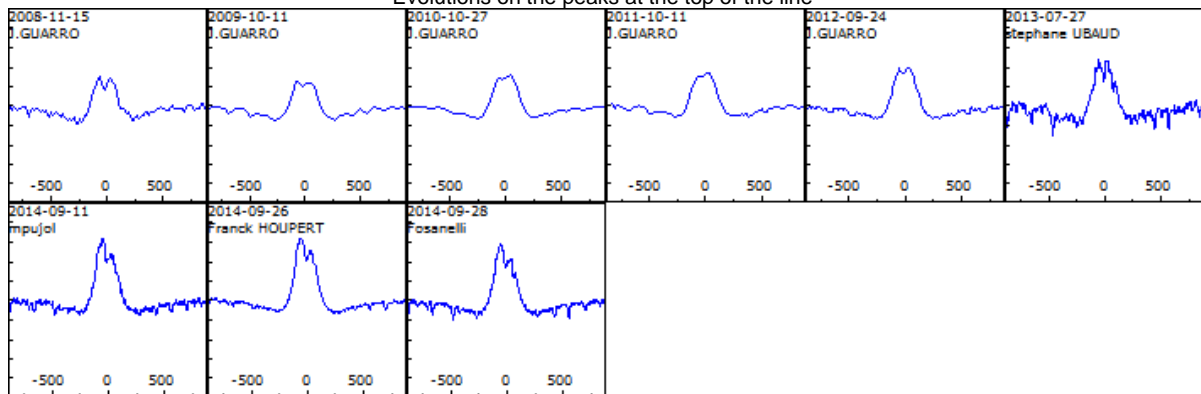
Evolution at the bottom of the line – only one year data





HD 201836

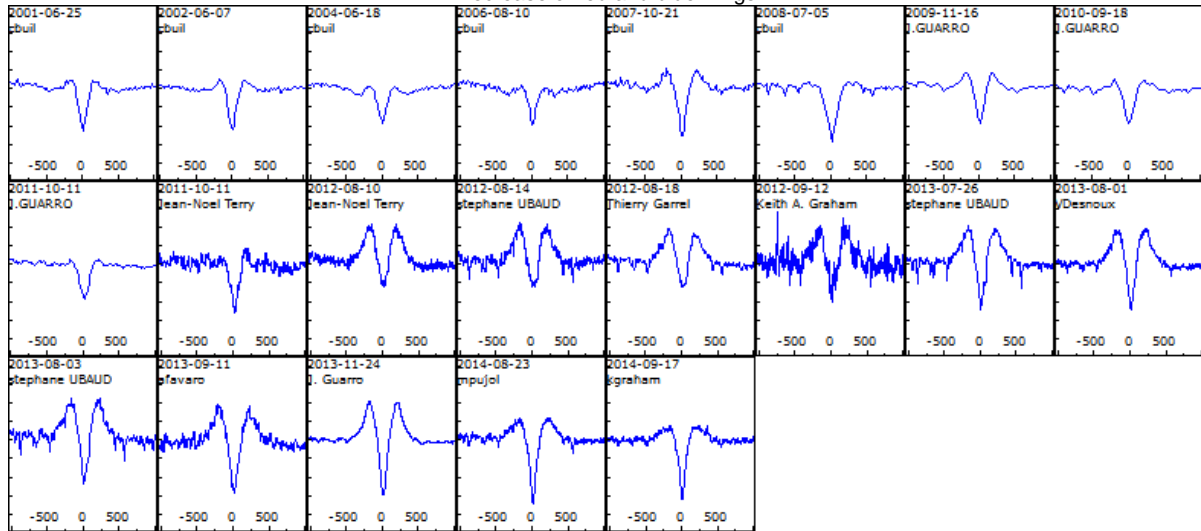
Evolutions on the peaks at the top of the line



Emission decrease of H-alpha line

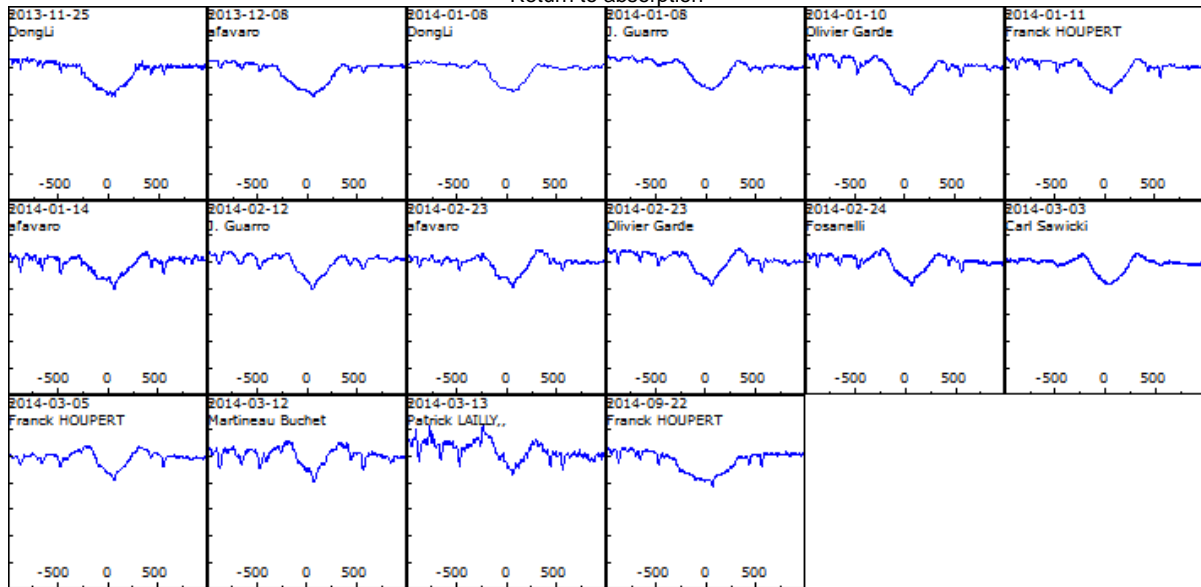
HD201733

Decrease of red and blue wings



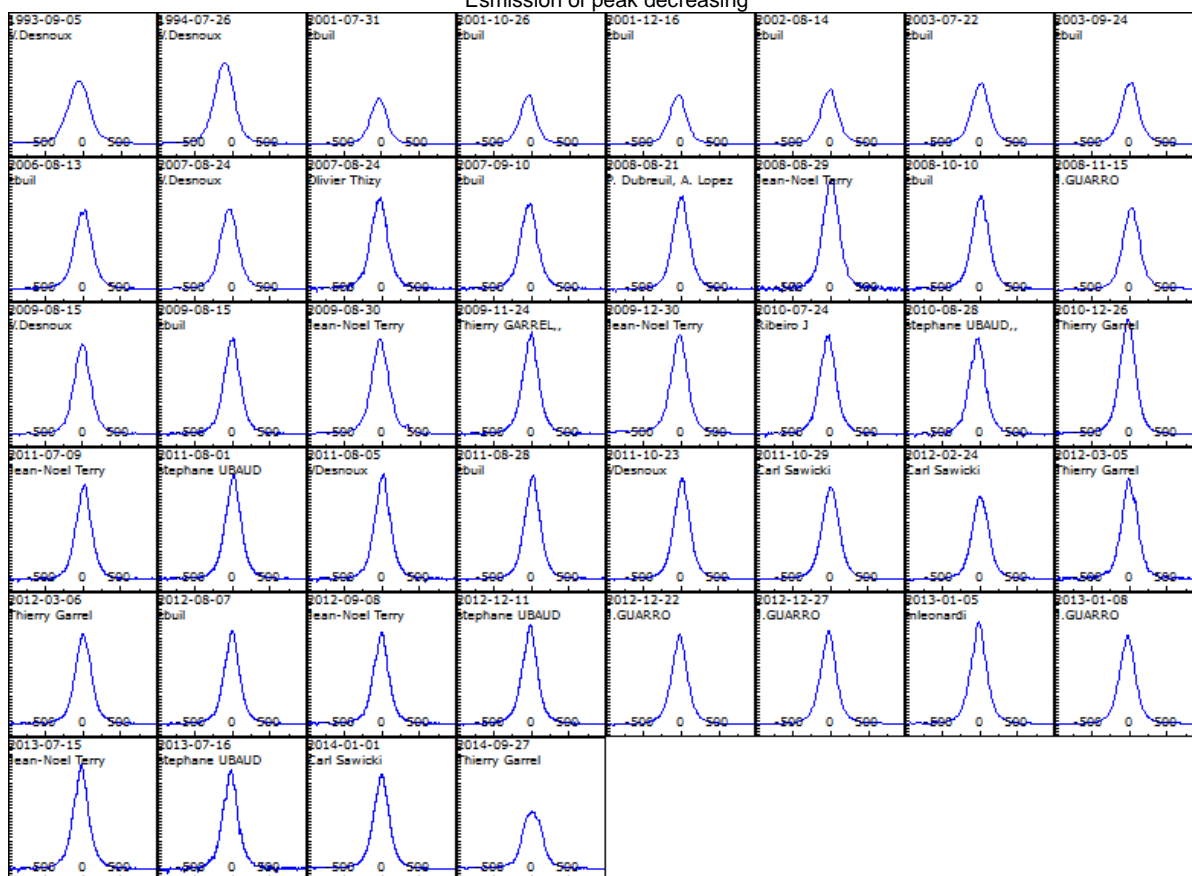
Lam Eri

Return to absorption



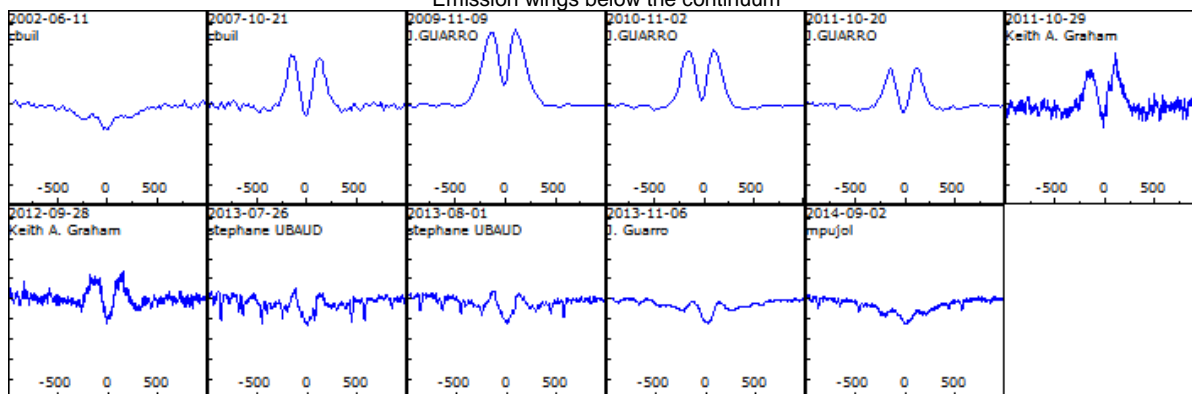
Omi Cas

Emission of peak decreasing



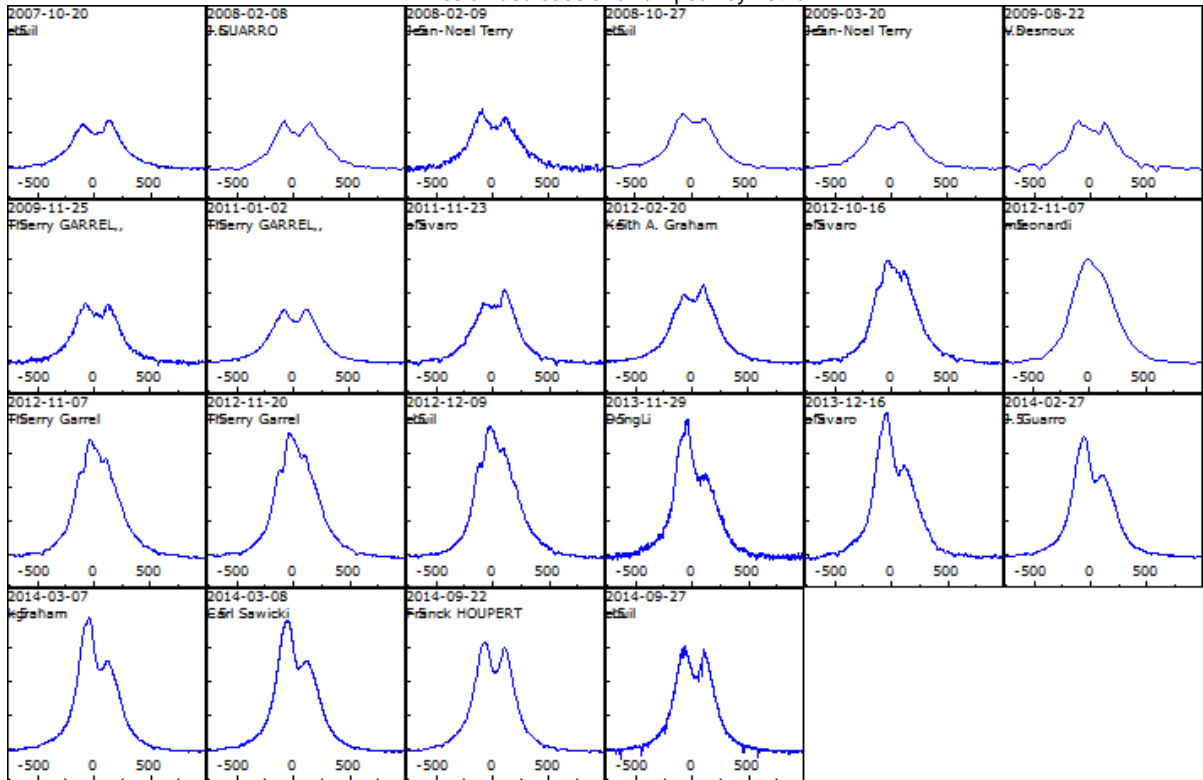
V2163 Cyg

Emission wings below the continuum



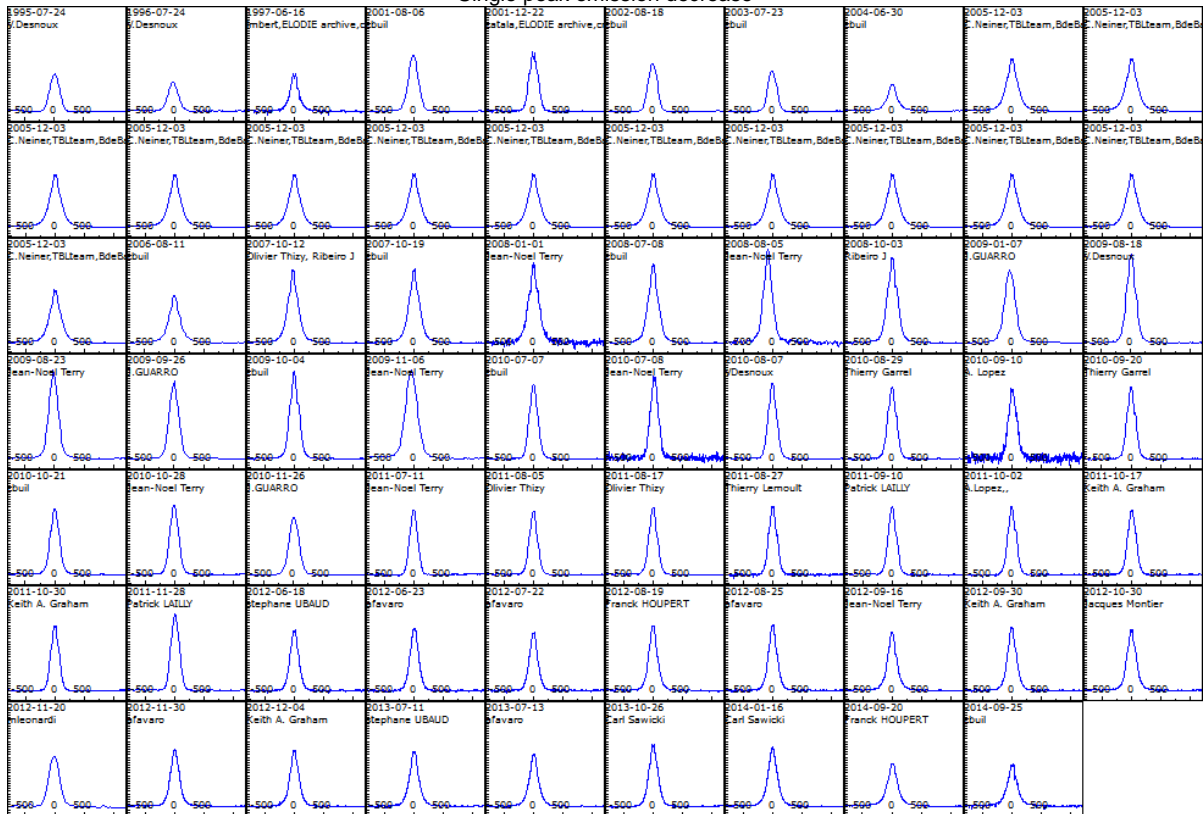
25 Ori

Emission decrease and V/R peak symmetric



31 Peg

Single peak emission decrease

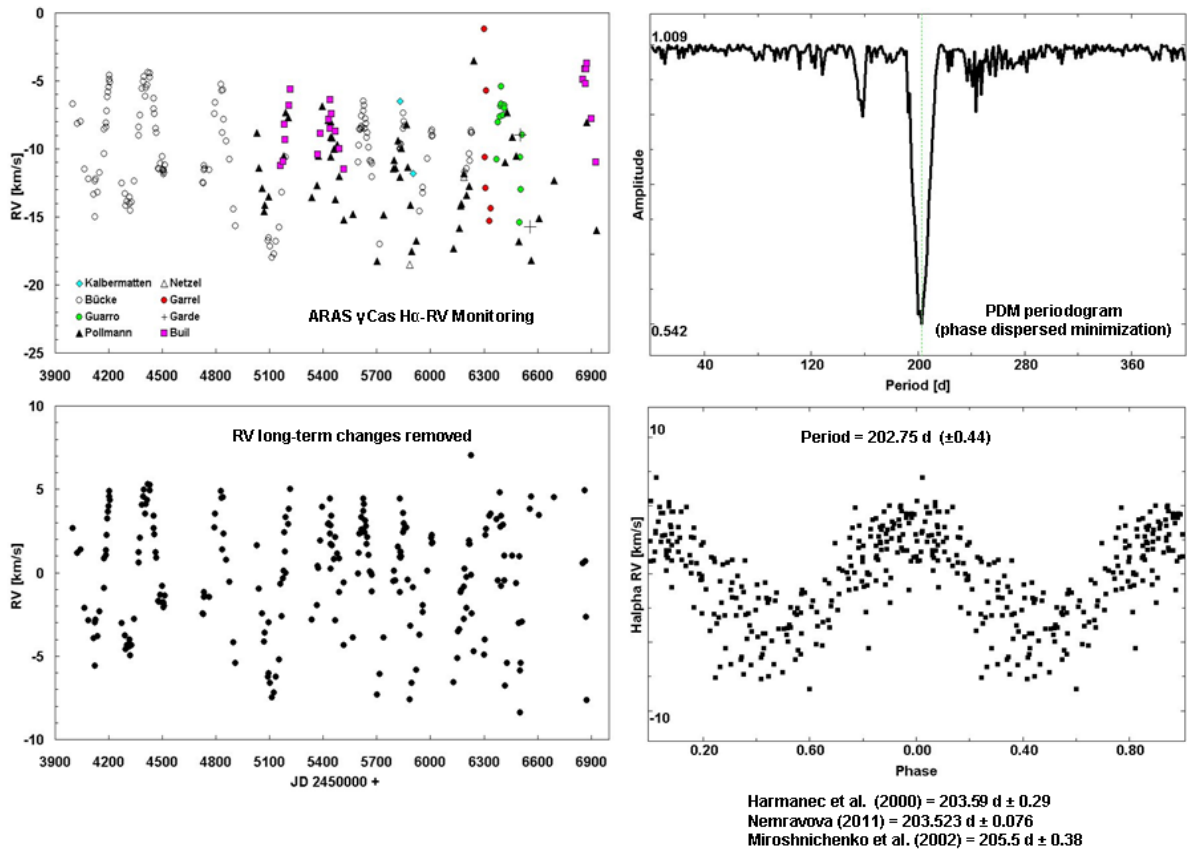


Be monitoring projects

By Ernst Pollmann

Halpna radial velocity of gamma Cas

The ARAS observation community is monitoring the Halpna radial velocity of gamma Cas since 09/2006 to 09/2014 (last communication to the group by me 2012/07/28). The meanwhile grown recordings (with very different spectrographs: LHIRES, LISA, VHRES) seems to be interesting to perform a new analysis (see the self explained attachment). Our period deviation compared with the periods of Nemravova, Miroshnichenko and Harmanec, results with high probability from use of the different spectrographs. Comments are welcome!

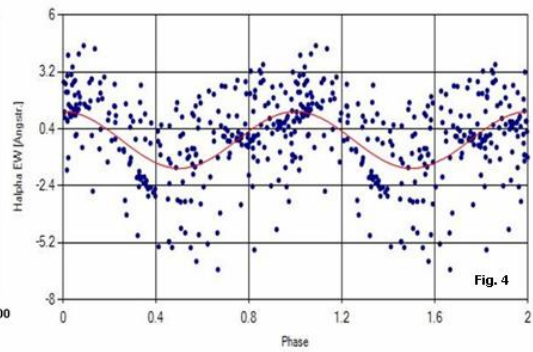
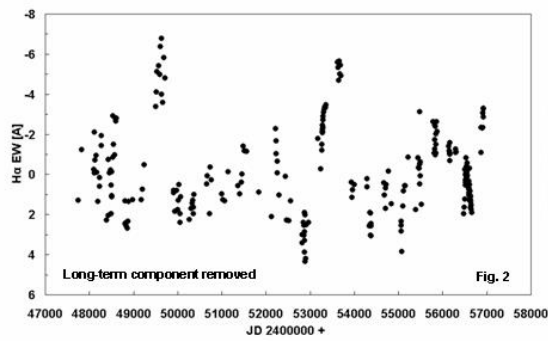
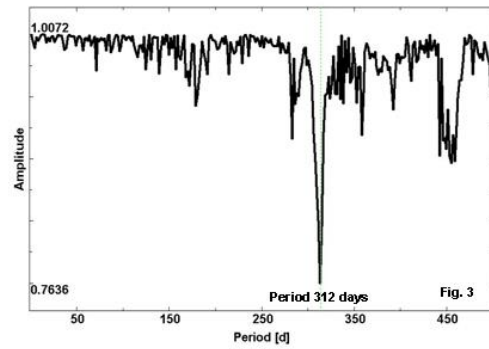
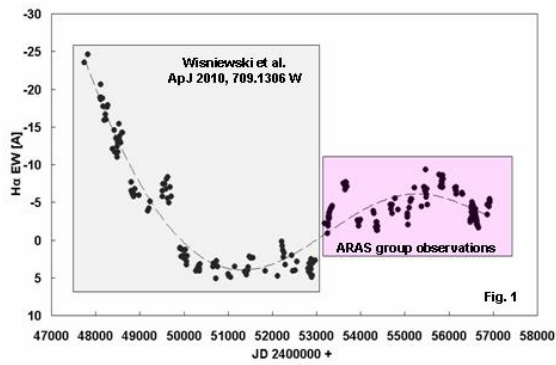


E. Pollmann, October 2014

Halpna EW Analysis of pi Aqr

Regarding the meaning of the "low" and "high" Halpna EW level in our pi Aqr paper (A&A, 560, A30, 2013), it seems to be interesting, to study the Halpna EW in its long-term behavior. That's why I want to show an analysis of a larger data set of Halpna EW (including data of Wisniewski et al., ApJ, 2010) from August 1989 until now (Fig. 1). Fig. 2 shows the Halpna EW long-term behavior with the removed long-term component. A PDM period analysis of the Fig. 2 data is shown in Fig. 3. In spite of the clear and very narrow period at 312 days, the phase diagram in Fig. 4 don't show a significant periodic behavior. The rms value (1.84 Angstr.) is clearly larger than the amplitude (1.40 Angstr.). It is

questionable, whether still more EW's lead to a higher clarity.



Authors

Valérie Desnoux

Valerie.desnoux@free.fr

Aras Site at <http://www.astrosurf.com/aras/>

BeSS database at <http://basebe.obspm.fr/basebe/>

ArasBeAM portal at <http://arasbeam.free.fr/>

Ernst Pollmann

ernst-pollmann@t-online.de

International Working Group ASPA

Active Spectroscopy in Astronomy

<http://www.astrospectroscopy.de>

<http://www.astronomie.de/astronomische-fachgebiete>